## Globular Cluster Abundance Anomalies and the Massive Binary Polluter Scenario

## Dorottya Szécsi

Nicolas đonzallez-Jimenez
Norbert Langer

Binary systems, their evolution and environments 1-5. September 2014, Ulaan Baatar

## A grid of low metallicity single stars

Szécsi et al. 2014 in prep.


## A grid of low metallicity single stars

Szécsi et al. 2014 in prep.


Abundance anomalies observed in Galactic Clusters (GCs)

## Abundance anomalies observed in GCs



## Abundance anomalies observed in GCs



## Abundance anomalies observed in GCs



## Abundance anomalies observed in GCs



## Abundance anomalies observed in GCs



## Abundance anomalies observed in GCs



## Abundance anomalies observed in GCs



## Abundance anomalies observed in GCs

Mg - Al anticorrelation



## Abundance anomalies observed in GCs

- extreme \& intermediate pop: polluted by hydrogen burning side products


## Abundance anomalies observed in GCs

- extreme \& intermediate pop: polluted by hydrogen burning side products
- CNO-cycle, Ne-Na and Mg-Al chains


## Abundance anomalies observed in GCs

- extreme \& intermediate pop: polluted by hydrogen burning side products
- CNO-cycle, $\mathrm{Ne}-\mathrm{Na}$ and Mg -Al chains
- either two generations of stars or accretion onto pre-MS low mass stars (Bastian+ 2013)


## Abundance anomalies observed in GCs

- extreme \& intermediate pop: polluted by hydrogen burning side products
- CNO-cycle, $\mathrm{Ne}-\mathrm{Na}$ and Mg -Al chains
- either two generations of stars or accretion onto pre-MS low mass stars (Bastian+ 2013)
- need: astrophysical source that can pollute the ISM


## Abundance anomalies observed in GCs

- extreme \& intermediate pop: polluted by hydrogen burning side products
- CNO-cycle, $\mathrm{Ne}-\mathrm{Na}$ and Mg -Al chains
- either two generations of stars or accretion onto pre-MS low mass stars (Bastian+ 2013)
- need: astrophysical source that can pollute the ISM
- caveat: only products of CNO-cycle (e.g. not He-burning or SN ejecta)


## Abundance anomalies observed in GCs

- extreme \& intermediate pop: polluted by hydrogen burning side products
- CNO-cycle, $\mathrm{Ne}-\mathrm{Na}$ and Mg -Al chains
- either two generations of stars or accretion onto pre-MS low mass stars (Bastian+ 2013)
- need: astrophysical source that can pollute the ISM
- caveat: only products of CNO-cycle (e.g. not He-burning or SN ejecta)
- caveat: material stays inside the grav. potential of the cluster (e.g. not fast stellar wind)


## Abundance anomalies observed in GCs

- extreme \& intermediate pop: polluted by hydrogen burning side products
- CNO-cycle, Ne -Na and Mg -Al chains
- either two generations of stars or accretion onto pre-MS low mass stars (Bastian+ 2013)
- need: astrophysical source that can pollute the ISM
- caveat: only products of CNO-cycle (e.g. not He-burning or SN ejecta)
- caveat: material stays inside the grav. potential of the cluster (e.g. not fast stellar wind)
- AGB stars: hot bottom burning (Ventura+2001)


## Abundance anomalies observed in GCs

- extreme \& intermediate pop: polluted by hydrogen burning side products
- CNO-cycle, $\mathrm{Ne}-\mathrm{Na}$ and Mg -Al chains
- either two generations of stars or accretion onto pre-MS low mass stars (Bastian+ 2013)
- need: astrophysical source that can pollute the ISM
- caveat: only products of CNO-cycle (e.g. not He-burning or SN ejecta)
- caveat: material stays inside the grav. potential of the cluster (e.g. not fast stellar wind)
- AGB stars: hot bottom burning (Ventura+2001)
- fast rotating massive stars: close to break-up (Decressin+2007)


## Abundance anomalies observed in GCs

- extreme \& intermediate pop: polluted by hydrogen burning side products
- CNO-cycle, $\mathrm{Ne}-\mathrm{Na}$ and Mg -Al chains
- either two generations of stars or accretion onto pre-MS low mass stars (Bastian+ 2013)
- need: astrophysical source that can pollute the ISM
- caveat: only products of CNO-cycle (e.g. not He-burning or SN ejecta)
- caveat: material stays inside the grav. potential of the cluster (e.g. not fast stellar wind)
- AGB stars: hot bottom burning (Ventura+2001)
- fast rotating massive stars: close to break-up (Decressin+2007)
- supermassive stars $\left(10^{4} \mathrm{M}_{\odot}\right)$ : continuum-driven wind (Denissenkov+2014)


## Abundance anomalies observed in GCs

- extreme \& intermediate pop: polluted by hydrogen burning side products
- CNO-cycle, Ne -Na and Mg -Al chains
- either two generations of stars or accretion onto pre-MS low mass stars (Bastian+ 2013)
- need: astrophysical source that can pollute the ISM
- caveat: only products of CNO-cycle (e.g. not He-burning or SN ejecta)
- caveat: material stays inside the grav. potential of the cluster (e.g. not fast stellar wind)
- AGB stars: hot bottom burning (Ventura+2001)
- fast rotating massive stars: close to break-up (Decressin+2007)
- supermassive stars $\left(10^{4} \mathrm{M}_{\odot}\right)$ : continuum-driven wind (Denissenkov+2014)
- massive binaries: non-conservative mass transfer (De Mink+ 2009)

The massive binary polluter scenario

## The massive binary polluter scenario

- interacting binary system, non-conservative mass transfer


## The massive binary polluter scenario

- interacting binary system, non-conservative mass transfer
- observational evidence for binaries loosing large amount of mass (see de Mink+ 2007 for a review)


## The massive binary polluter scenario

- interacting binary system, non-conservative mass transfer
- observational evidence for binaries loosing large amount of mass (see de Mink+ 2007 for a review)
- after H-exhaustion: primary expands, secondary accretes $\rightarrow$ spins up


## The massive binary polluter scenario

- interacting binary system, non-conservative mass transfer
- observational evidence for binaries loosing large amount of mass (see de Mink+ 2007 for a review)
- after H-exhaustion: primary expands, secondary accretes $\rightarrow$ spins up
- reaches critical rotation $\rightarrow$ mass is ejected from the system
- slow ejecta


## The massive binary polluter scenario

- interacting binary system, non-conservative mass transfer
- observational evidence for binaries loosing large amount of mass (see de Mink+ 2007 for a review)
- after H-exhaustion: primary expands, secondary accretes $\rightarrow$ spins up
- reaches critical rotation $\rightarrow$ mass is ejected from the system
- slow ejecta
- deeper layers of primary envelope: nuclearly processed material!
- hydrogen burning products


## The massive binary polluter scenario

- interacting binary system, non-conservative mass transfer
- observational evidence for binaries loosing large amount of mass (see de Mink+ 2007 for a review)
- after H-exhaustion: primary expands, secondary accretes $\rightarrow$ spins up
- reaches critical rotation $\rightarrow$ mass is ejected from the system
- slow ejecta
- deeper layers of primary envelope: nuclearly processed material!
- hydrogen burning products
- De Mink+ 2009: $20 \mathrm{M}_{\odot}+15 \mathrm{M}_{\odot}+12$ days $\left(\sim 0.025 \mathrm{Z}_{\odot}\right)$


## The massive binary polluter scenario




The massive binary polluter scenario

## The massive binary polluter scenario




## The massive binary polluter scenario

$$
\mathrm{m} 1=020 \mathrm{M}_{\odot} \quad \mathrm{m} 2=015 \mathrm{M}_{\odot} \quad \mathrm{p}=12.00 \mathrm{~d}
$$

de Mink et al. 2009


## Other systems?

More massive systems...

## Other systems?

More massive systems...

- How much mass would they possibly eject?


## Other systems?

More massive systems...

- How much mass would they possibly eject?
- Would their ejecta composition reproduce the observed anticorrelations?


## Other systems?

More massive systems...

- How much mass would they possibly eject?
- Would their ejecta composition reproduce the observed anticorrelations?

A grid of low metallicity single stars...

## A grid of low metallicity single stars

Szécsi et al. 2014 in prep.


## A grid of low metallicity single stars



Single star approach
to the Massive Binary Polluter Scenario

## Single star approach

Szécsi et al. 2014 in prep.


## Single star approach

Szécsi et al. 2014 in prep.


## Single star approach

Szécsi et al. 2014 in prep.


## Single star approach

Szécsi et al. 2014 in prep.


## Single star approach

Szécsi et al. 2014 in prep.


## Single star approach

Szécsi et al. 2014 in prep.


## Single star approach

Szécsi et al. 2014 in prep.


## Composition and size of primary envelope



## Composition and size of primary envelope



## Composition and size of primary envelope



## Single star approach

## Assumptions:

## Single star approach

## Assumptions:

- $\tau_{\text {masstransfer }}<\tau_{\text {mixing }}, \tau_{\text {nucl }}$


## Single star approach

Assumptions:

- $\tau_{\text {masstransfer }}<\tau_{\text {mixing }}, \tau_{n u c l}$
- orbit is not (fully) synchronized


## Single star approach

Assumptions:

- $\tau_{\text {masstransfer }}<\tau_{\text {mixing }}, \tau_{\text {nucl }}$
- orbit is not (fully) synchronized
- whole envelope is ejected


## Single star approach

Assumptions:

- $\tau_{\text {masstransfer }}<\tau_{\text {mixing }}, \tau_{n u c l}$
- orbit is not (fully) synchronized
- whole envelope is ejected
$\rightarrow$ detailed binary simulations still needed


## Single star approach

Assumptions:

- $\tau_{\text {masstransfer }}<\tau_{\text {mixing }}, \tau_{n u c l}$
- orbit is not (fully) synchronized
- whole envelope is ejected
$\rightarrow$ detailed binary simulations still needed
Advantages:


## Single star approach

Assumptions:

- $\tau_{\text {masstransfer }}<\tau_{\text {mixing }}, \tau_{n u c l}$
- orbit is not (fully) synchronized
- whole envelope is ejected
$\rightarrow$ detailed binary simulations still needed
Advantages:
- detailed calculations of single stars are less difficult $\rightarrow$ cover a broad parameter space


## Single star approach

Assumptions:

- $\tau_{\text {masstransfer }}<\tau_{\text {mixing }}, \tau_{n u c l}$
- orbit is not (fully) synchronized
- whole envelope is ejected
$\rightarrow$ detailed binary simulations still needed
Advantages:
- detailed calculations of single stars are less difficult $\rightarrow$ cover a broad parameter space
- in case of simulating binaries: it helps to decide which masses, mass ratios and periods to simulate and what to expect


## Single star approach

Assumptions:

- $\tau_{\text {masstransfer }}<\tau_{\text {mixing }}, \tau_{n u c l}$
- orbit is not (fully) synchronized
- whole envelope is ejected
$\rightarrow$ detailed binary simulations still needed
Advantages:
- detailed calculations of single stars are less difficult $\rightarrow$ cover a broad parameter space
- in case of simulating binaries: it helps to decide which masses, mass ratios and periods to simulate and what to expect
- give constraints on the massive binary polluter scenaro even without detailed binary simulations


## Single star approach: size of primary envelope



## Single star approach: size of primary envelope



## Single star approach: size of primary envelope



## Single star approach: size of primary envelope



## Single star approach: size of primary envelope



Compared to observations:
O - Na anticorrelation

## Compared to observations: O - Na anticorr.



## Compared to observations: O - Na anticorr.



## Compared to observations: O - Na anticorr.



## Compared to observations: $\mathrm{O}-\mathrm{Na}$ anticorr.



## Compared to observations:

 Mg - Al anticorrelationCompared to observations: $\mathrm{Mg}-\mathrm{Al}$ anticorr.


Compared to observations: $\mathrm{Mg}-\mathrm{Al}$ anticorr.


Compared to observations: $\mathrm{Mg}-\mathrm{Al}$ anticorr.


## Compared to observations: Mg - Al anticorr.



## Compared to observations: Mg - Al anticorr.



## Summary of the results

- Mg-Al problem - possible solutions:


## Summary of the results

- Mg-Al problem - possible solutions:
- $\sim \mathbf{1 0 0} \mathbf{M}_{\odot}$ primary losing H -shell burning products


## Summary of the results

- Mg-Al problem - possible solutions:
- $\sim \mathbf{1 0 0} \mathbf{M}_{\odot}$ primary losing H -shell burning products and/or


## Summary of the results

- Mg-Al problem - possible solutions:
- $\sim \mathbf{1 0 0} \mathbf{M}_{\odot}$ primary losing H -shell burning products and/or
- the presence of $\gtrsim \mathbf{5 0 0} \mathbf{M}_{\odot}$ primaries in the cluster


## Summary of the results

- Mg-Al problem - possible solutions:
- $\sim \mathbf{1 0 0} \mathbf{M}_{\odot}$ primary losing H-shell burning products and/or
- the presence of $\gtrsim 500 \mathbf{M}_{\odot}$ primaries in the cluster
- Envelope mass as a function of primary mass:


## Summary of the results

- Mg-Al problem - possible solutions:
- $\sim \mathbf{1 0 0} \mathbf{M}_{\odot}$ primary losing H -shell burning products and/or
- the presence of $\gtrsim 500 \mathbf{M}_{\odot}$ primaries in the cluster
- Envelope mass as a function of primary mass:
- assumption of de Mink+09 and Pranczos+Charbonnel'06 is supported by my calculations of single stars


## Summary of the results

- Mg-Al problem - possible solutions:
- $\sim \mathbf{1 0 0} \mathbf{M}_{\odot}$ primary losing H-shell burning products and/or
- the presence of $\gtrsim \mathbf{5 0 0} \mathbf{M}_{\odot}$ primaries in the cluster
- Envelope mass as a function of primary mass:
- assumption of de Mink+09 and Pranczos+Charbonnel'06 is supported by my calculations of single stars
- extended for higher masses (up to $\sim 575 \mathrm{M}_{\odot}$ )


## Final notes

Work in progress, first steps presented.

## Final notes

Work in progress, first steps presented. Aims of this talk:

## Final notes

Work in progress, first steps presented. Aims of this talk:

- present GC abundance anom. \& massive binary polluters


## Final notes

Work in progress, first steps presented. Aims of this talk:

- present GC abundance anom. \& massive binary polluters
- present the idea of single star approach


## Final notes

Work in progress, first steps presented. Aims of this talk:

- present GC abundance anom. \& massive binary polluters
- present the idea of single star approach
- present my grid of low metallicity single stars


## Final notes

Work in progress, first steps presented. Aims of this talk:

- present GC abundance anom. \& massive binary polluters
- present the idea of single star approach
- present my grid of low metallicity single stars
- broad range of masses and rotations


## Final notes

Work in progress, first steps presented. Aims of this talk:

- present GC abundance anom. \& massive binary polluters
- present the idea of single star approach
- present my grid of low metallicity single stars
- broad range of masses and rotations
- composition suitable for GCs, dwarf galaxies, high z objects


## Final notes

Work in progress, first steps presented. Aims of this talk:

- present GC abundance anom. \& massive binary polluters
- present the idea of single star approach
- present my grid of low metallicity single stars
- broad range of masses and rotations
- composition suitable for GCs, dwarf galaxies, high z objects
- give constraints on the envelope mass, test assumptions of de Mink+09, Bastian+13


## Final notes

Work in progress, first steps presented. Aims of this talk:

- present GC abundance anom. \& massive binary polluters
- present the idea of single star approach
- present my grid of low metallicity single stars
- broad range of masses and rotations
- composition suitable for GCs, dwarf galaxies, high z objects
- give constraints on the envelope mass, test assumptions of de Mink+09, Bastian+13 $\checkmark$


## Final notes

Work in progress, first steps presented. Aims of this talk:

- present GC abundance anom. \& massive binary polluters
- present the idea of single star approach
- present my grid of low metallicity single stars
- broad range of masses and rotations
- composition suitable for GCs, dwarf galaxies, high z objects
- give constraints on the envelope mass, test assumptions of de Mink+09, Bastian+13 $\checkmark$
- test if higher masses / wider periods could help the Mg-Al problem


## Final notes

Work in progress, first steps presented. Aims of this talk:

- present GC abundance anom. \& massive binary polluters
- present the idea of single star approach
- present my grid of low metallicity single stars
- broad range of masses and rotations
- composition suitable for GCs, dwarf galaxies, high z objects
- give constraints on the envelope mass, test assumptions of de Mink+09, Bastian+13 $\checkmark$
- test if higher masses / wider periods could help the Mg-Al problem $\checkmark$


## Final notes

Work in progress, first steps presented. Aims of this talk:

- present GC abundance anom. \& massive binary polluters
- present the idea of single star approach
- present my grid of low metallicity single stars
- broad range of masses and rotations
- composition suitable for GCs, dwarf galaxies, high z objects
- give constraints on the envelope mass, test assumptions of de Mink+09, Bastian+13 $\checkmark$
- test if higher masses / wider periods could help the Mg-Al problem $\checkmark$
Still need to investigate / question marks:
- mass budget; effects of rotation
- $\mathrm{Al}-\mathrm{Mg}$ chain: update the reaction rates
- systems interacting at a young age - binary simulations needed


## Final notes

Work in progress, first steps presented. Aims of this talk:

- present GC abundance anom. \& massive binary polluters
- present the idea of single star approach
- present my grid of low metallicity single stars
- broad range of masses and rotations
- composition suitable for GCs, dwarf galaxies, high z objects
- give constraints on the envelope mass, test assumptions of de Mink+09, Bastian+13 $\checkmark$
- test if higher masses / wider periods could help the Mg-Al problem $\checkmark$
Still need to investigate / question marks:
- mass budget; effects of rotation
- $\mathrm{Al}-\mathrm{Mg}$ chain: update the reaction rates
- systems interacting at a young age - binary simulations needed

Open to suggestions, comments and questions!

## Final notes

Work in progress, first steps presented. Aims of this talk:

- present GC abundance anom. \& massive binary polluters
- present the idea of single star approach
- present my grid of low metallicity single stars
- broad range of masses and rotations
- composition suitable for GCs, dwarf galaxies, high z objects
- give constraints on the envelope mass, test assumptions of de Mink+09, Bastian+13 $\checkmark$
- test if higher masses / wider periods could help the Mg-Al problem $\checkmark$
Still need to investigate / question marks:
Thank you for your attention!
- mass budget; effects of rotation
- $\mathrm{Al}-\mathrm{Mg}$ chain: update the reaction rates
- systems interacting at a young age - binary simulations needed

Open to suggestions, comments and questions!

